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Search for dark matter signals with the H.E.S.S. Inner Galaxy Survey

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OUTLINE

- Introduction: indirect Dark Matter search in gamma rays
- Inner Galaxy Survey performed by H.E.S.S.
- H.E.S.S. data analysis
- Computation of the upper limits on $\langle\sigma v\rangle$
- Conclusions

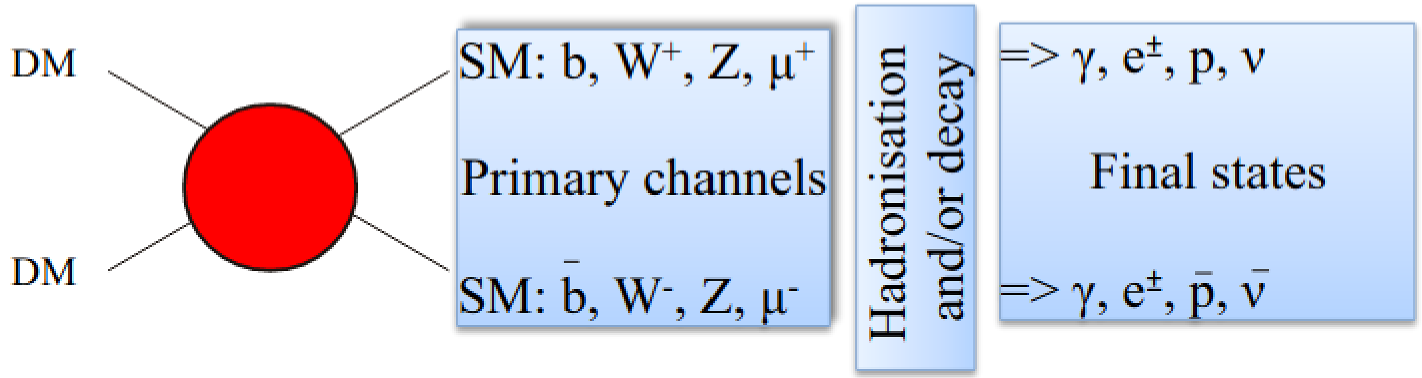
Indirect Dark Matter search

- Growing astrophysical and cosmological evidence about the existence of Dark Matter (DM).

- **WIMPs → one of the most compelling DM particle candidates.**

- WIMPs created thermally in the Early Universe:
  - Annihilation cross section expected for thermal WIMPs ($\langle\sigma v\rangle_{th} = 3x10^{-26}\ \mathrm{cm}^3\ \mathrm{s}^{-1}$).

- WIMPs can self-annihilate and produce Standard Model particles in the final states eventually detectable by satellite (*Fermi*-LAT) and ground-based experiments (H.E.S.S., VERITAS).



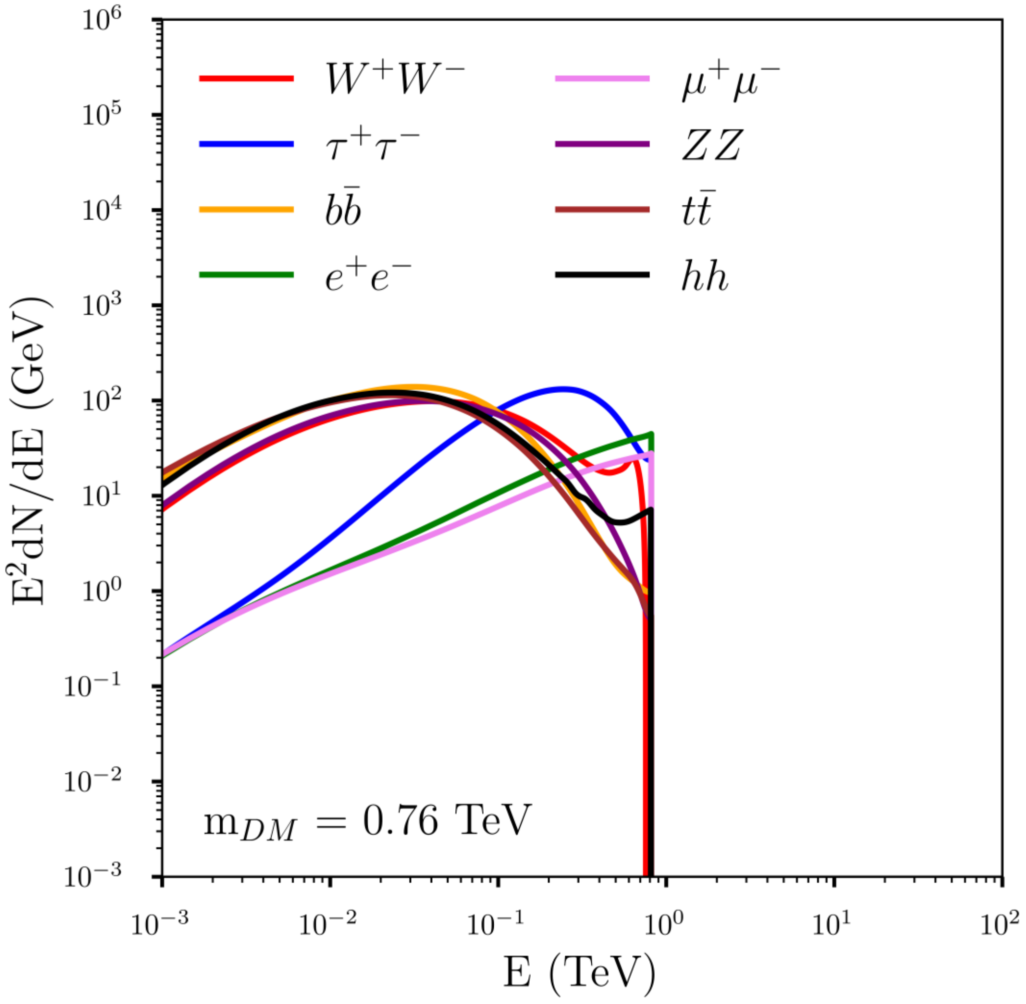
Indirect Dark Matter search in gamma rays

- WIMPs can self-annihilate and produce gamma-rays eventually detectable by H.E.S.S.

- Assuming annihilation process almost at rest:
  - A smoking-gun signature for DM is a very distinct energy cut-off, close to the DM particle mass.

- **Gamma-ray flux expected from DM annihilations:**

$$\frac{d\phi_\gamma}{dE}(E_\gamma, \Delta\Omega) = \frac{\langle\sigma v\rangle J(\Delta\Omega)}{8\pi m_{DM}^2} \sum_f Br_f \frac{dN_f}{dE_\gamma}$$



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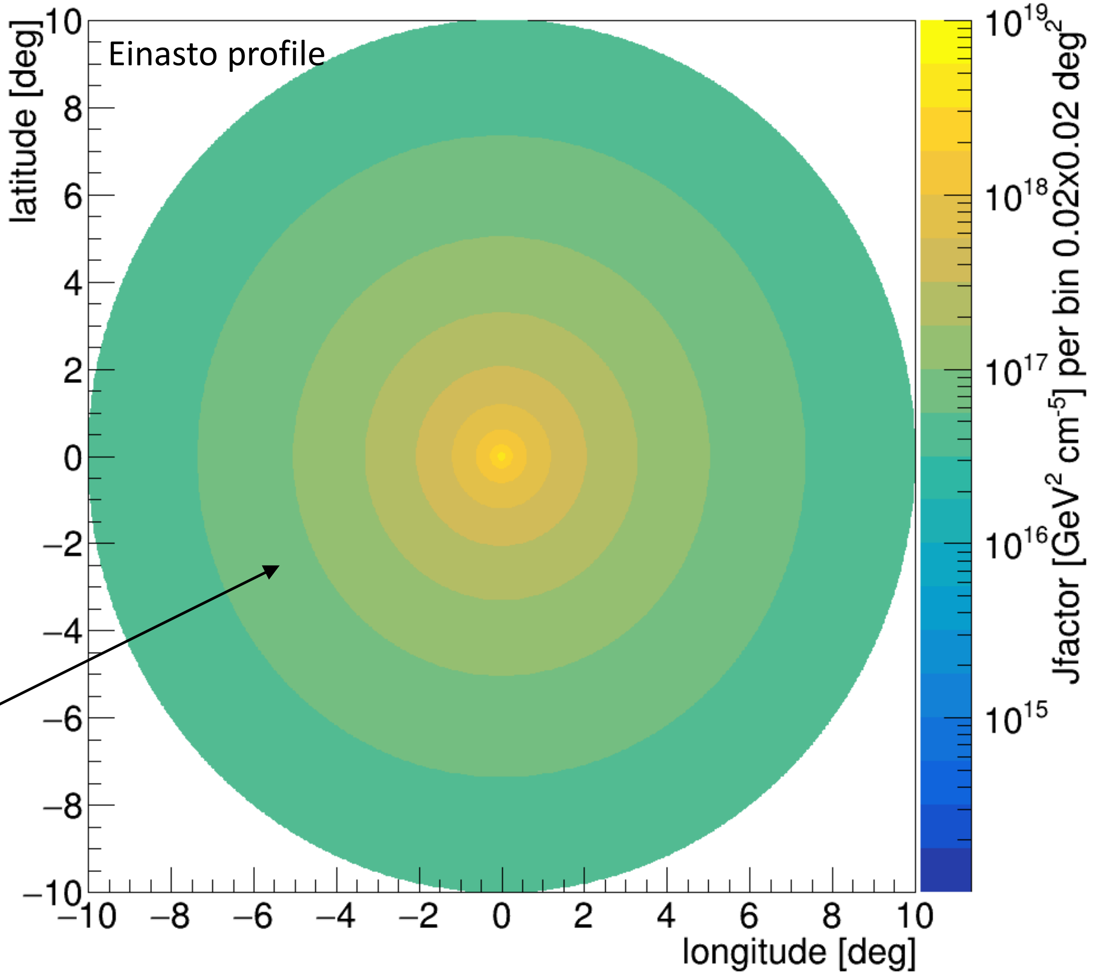
- **Astrophysical term $J(\Delta\Omega) = \int \rho^2\,(r(s,\theta))\,ds\,d\Omega$:**
  - Model needed for the density profile;
  - Dependence on dark matter halo modeling.



Indirect Dark Matter targets in gamma rays

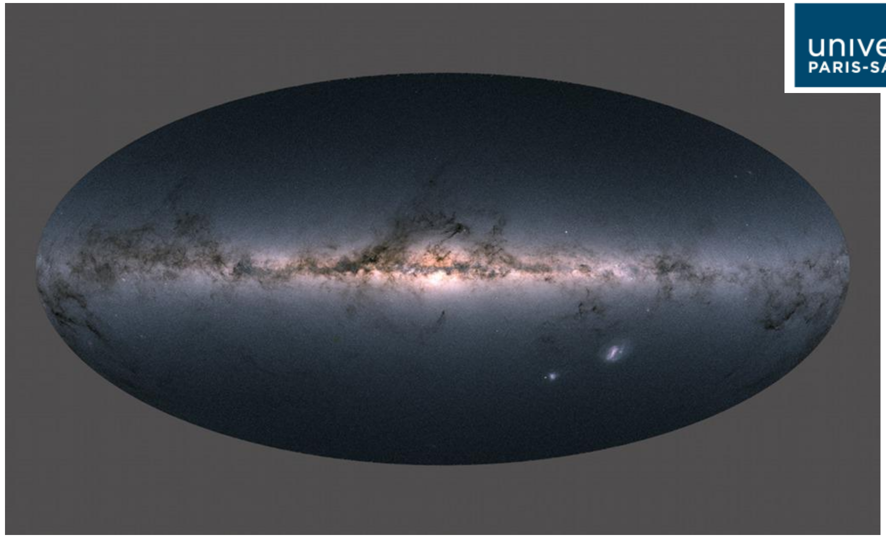
- Astrophysical term $J(\Delta\Omega) = \int \rho^2\,(r(s,\theta))ds d\Omega$:
  - Model needed for the density profile;
  - Dependence on dark matter halo modeling.

- **Most promising candidates for DM detection:**
  - **Galactic Center region** and nearby dwarf galaxies.

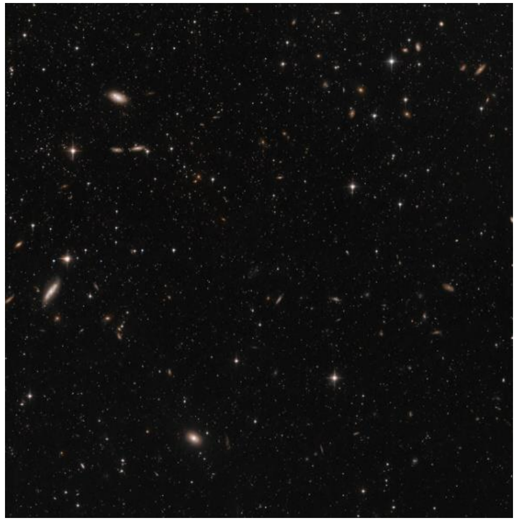
  Refs. Abdalla et al. [H.E.S.S. collaboration], Phys. Rev. Lett. 2016
  Abdalla et al. [H.E.S.S. collaboration], Phys. Rev. Lett. 2018
  Abdalla et al. [H.E.S.S. collaboration], JCAP 2018
  Abdalla et al. [H.E.S.S. collaboration], Phys. Rev. D. 2020

- Other compelling and complementary DM targets are:
  - Dark Matter subhalos populating the Galactic Halo.

  Ref. Abdalla et al. [H.E.S.S. collaboration], 2021, arXiv: 2106.00551



Grand all sky of our Milky Way and nearby galaxies. Credit: Gaia's satellite second data release



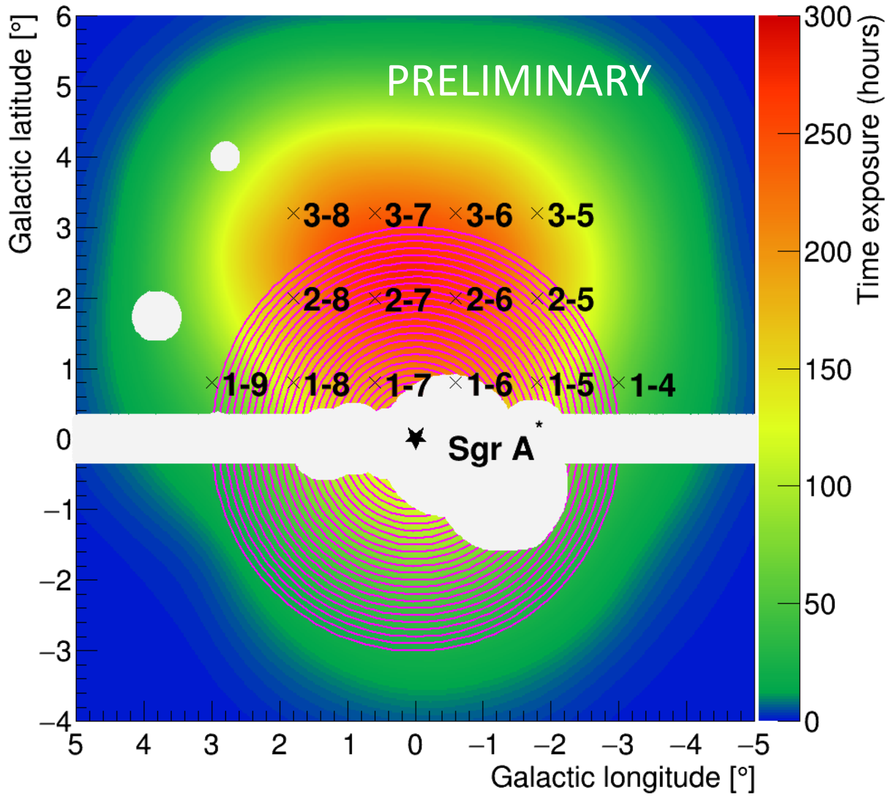
A small portion of the Sculptor dwarf galaxy, a satellite galaxy of the Milky Way, as observed by the NASA/ESA Hubble Space Telescope. This image shows one of two different pointings of the telescope.

Inner Galaxy Survey (IGS) performed by H.E.S.S.

- **The first ever conducted VHE gamma-ray survey of the Galactic Center (GC) region.**

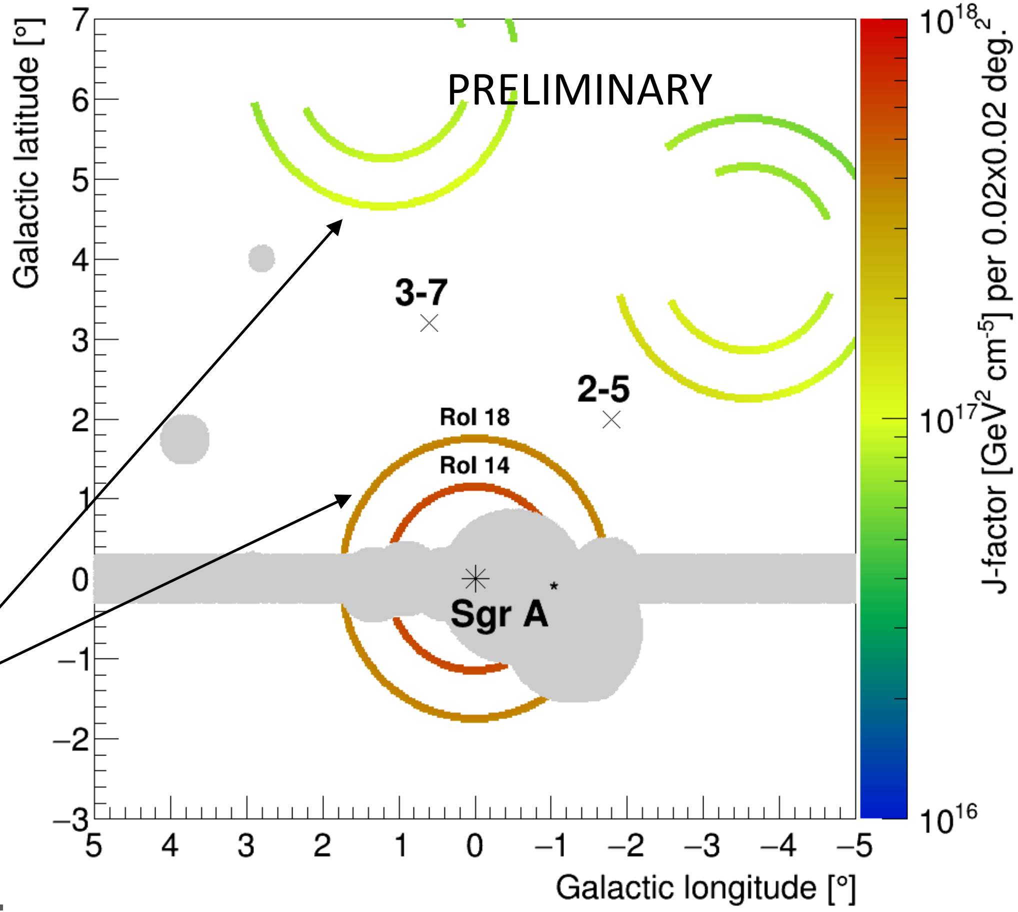
➢ Aim: to provide **unprecedented sensitivity** to DM signals in the GC region.

- Dataset: 2014-2020 observations of the GC region.
- 2014-2020 exposure map with IGS pointing positions:
  - Exposure up to b ≈ 6°;
  - 25 regions of interest (ROI) defined to search for DM: 0.1°-width open rings;
  - Set of exclusion regions to avoid gamma-ray contamination in the ROIs.



Background measurement, ON/OFF construction

- Definition of the ON region: 25 ROI.
- **Reflected background method:**
  - OFF region:
    - Symmetric to the ON region wrt the pointing position
      - Same FoV and acceptance;
    - The excluded regions are cut symmetrically
      - Same solid angle size;
    - Cut overlapping areas and areas where OFF is closer to GC than the ON:
      - The DM signal in the ON region is always higher than in the OFF region.



- Repeated for all the 25 ROI and over the ~1300 runs.

Likelihood analysis method and Test Statistics definition

- **2D binned Poisson likelihood function** exploits spatial and spectral DM features: bins in energy (i) and space (j):

$$\mathcal{L}_{i,j}(N_{S,ij}, N_{B,ij}, \beta_{ij} | N_{ON,ij}, N_{OFF,ij}, \alpha_j) = \frac{[\beta_{ij}(N_{S,ij} + N_{B,ij})]^{N_{ON,ij}}}{N_{ON,ij}!} e^{-\beta_{ij}(N_{S,ij}+N_{B,ij})} \frac{[\beta_{ij}(N'_{S,ij} + \alpha_j N_{B,ij})]^{N_{OFF,ij}}}{N_{OFF,ij}!} e^{-\beta_{ij}(N'_{S,ij}+\alpha_j N_{B,ij})} e^{\frac{(1-\beta_{ij})^2}{\sigma_\beta}}$$

- Total likelihood function: $\mathcal{L} = \prod_{i,j} \mathcal{L}_{i,j}$

- $N_{ON,ij}$ and $N_{OFF,ij}$ $\rightarrow$ number of measured events in spatial ON and OFF regions;
- $N_{S,ij} + N_{B,ij}$ $\rightarrow$ expected total number of events in the spatial ON region;
- $N'_{S,ij} + \alpha_j N_{B,ij}$ $\rightarrow$ expected total number of events in the spatial OFF region;
- $\alpha_j = {}^{\Delta\Omega_{ON}}/_{\Delta\Omega_{OFF}}$ $\rightarrow$ ratio between angular size of ON and OFF regions.

Likelihood analysis method and Test Statistics definition

- 2D binned Poisson likelihood function exploits spatial and spectral DM features: bins in energy (i) and space (j):

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- Total likelihood function: $\mathcal{L} = \prod \mathcal{L}_{i,j}$

- **The systematic uncertainties can be included via a nuisance parameter;**

  Refs: Silverwood, et al, JCAP03, 055 (2015); Lefranc, et al. Phys. Rev. D91, 122003 (2015); CTA DM Programme (2019)

  - A value of 1% is used for the determination of the limits: σ_β=0.01
  - The value of β is determined via conditional maximization
    - β is computed for each energy and spatial bins, i.e., $\beta_{i,j}$.

Likelihood analysis method and Test Statistics definition

- 2D binned Poisson likelihood function exploits spatial and spectral DM features:
  bins in energy (i) and space (j):

$$\mathcal{L}_{i,j}(N_{S,ij}, N_{B,ij}, \beta_{ij} | N_{ON,ij}, N_{OFF,ij}, \alpha_j) = \frac{[\beta_{ij}(N_{S,ij} + N_{B,ij})]^{N_{ON,ij}}}{N_{ON,ij}!} e^{-\beta_{ij}(N_{S,ij}+N_{B,ij})} \frac{[\beta_{ij}(N'_{S,ij} + \alpha_j N_{B,ij})]^{N_{OFF,ij}}}{N_{OFF,ij}!} e^{-\beta_{ij}(N'_{S,ij}+\alpha_j N_{B,ij})} e^{\frac{(1-\beta_{ij})^2}{\sigma_\beta}}$$

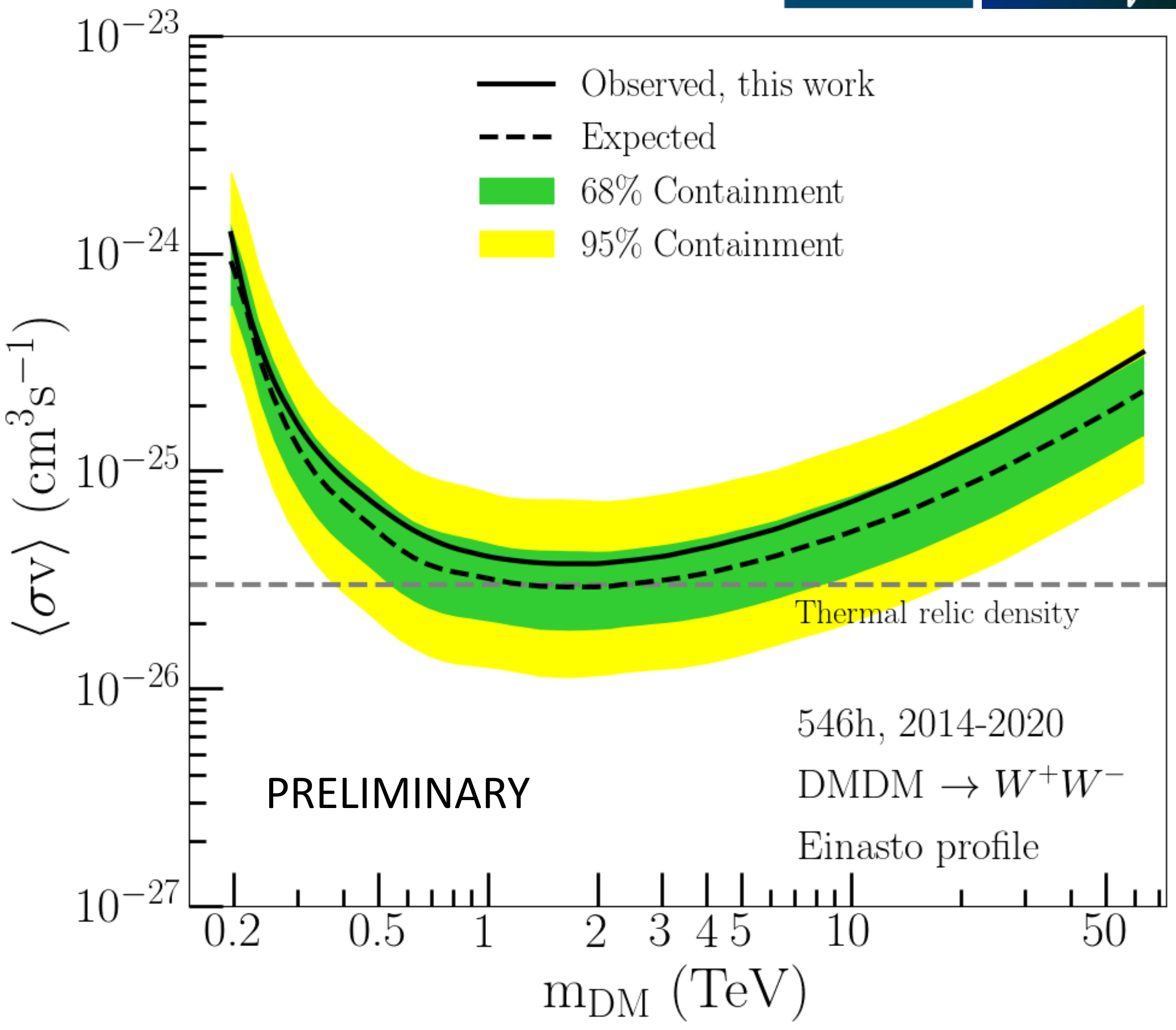
- Total likelihood function: $\mathcal{L} = \prod \mathcal{L}_{i,j}$

- In absence of any significant excess in the FoV:
  → **95% C.L. upper limits on the free parameter <σv> from a log-likelihood ratio test statistics (TS).**
  Ref. Cowan, G., Cranmer, K., Gross, E. *et al. Eur. Phys. J. C* **71**, 1554 (2011)

- **Computation of expected limits and containment bands:**
  - Independent Poisson realizations for the ON and OFF measurements;
    → mean and std deviation derived from the distribution of the obtained <σv> values.

Current expected and observed upper limits

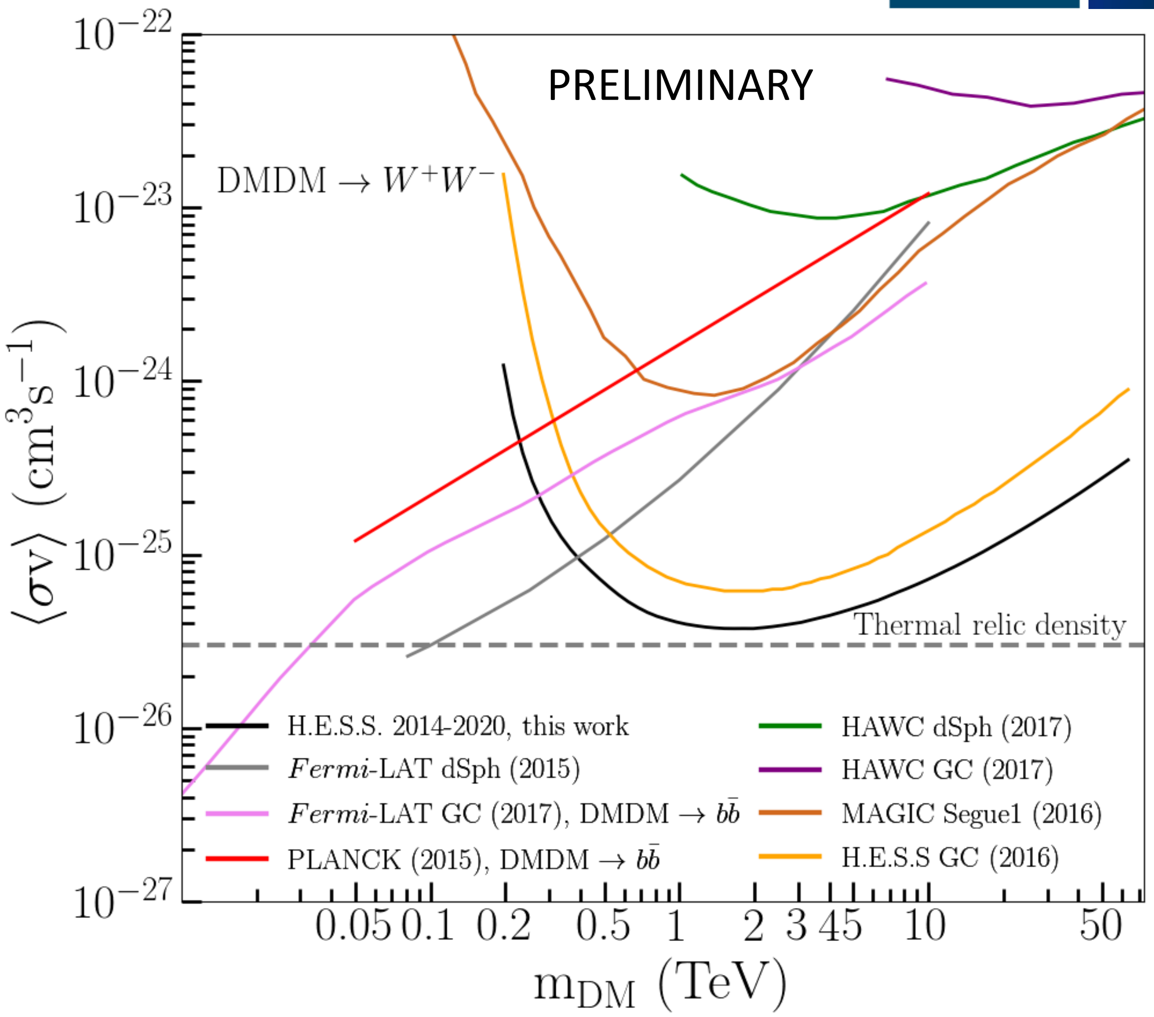
- No significant excess in the FoV:
  → 95% C.L. upper limits on <σv> from the TS;
- **H.E.S.S. upper limits;**
- Independent Poisson realizations for N_{ON} and N_{OFF} in the computation of the expected limits;
- Containment bands plotted at 1σ and 2σ level;
- Systematic uncertainty included in the limits via a nuisance parameter in the likelihood function.



Current observed upper limits

- H.E.S.S. upper limits.
- *Fermi*-LAT dSph and GC, HAWC dSph and GC, MAGIC Segue 1, PLANCK CMB, H.E.S.S. GC (2016) and this work.

→ Most constraining limits in the TeV-energy range.



Conclusions

- IGS campaign with pointing positions up to 3.2° is very fruitful:
    - Around 546 hours of high-quality data from 2014 to 2020.
- Computation of 95% C.L. expected and observed upper limits including systematic uncertainty.
- VHE observations of the GC region are unique for the study of the WIMP paradigm.
- With the unprecedented IGS dataset:
  → strongest constraints obtained in the TeV mass range.
- Limits are computed in other channels → can challenge the thermal relic scale.
- The IGS is one of the legacy of the H.E.S.S. collaboration and it paves the way for CTA.